

Hot Jupiters: Lands of Plenty

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Abstract. In late August 2005, 80 researchers from more than 15 countries convened for a 4-day conference entitled “The Tenth Anniversary of 51 Peg b: Status and Prospects for Hot Jupiter Studies”. The meeting was held at l’Observatoire de Haute-Provence, the location of the 1.93-m telescope and ELODIE spectrograph used to discover the planetary companion to 51 Peg roughly 10 years ago. I summarize several dominant themes that emerged from the meeting, including (i) recent improvements in the precision of radial velocity measurements of nearby, Sun-like stars, (ii) the continued value of individual, newly-discovered planets of novel character to expand the parameter space with which the theory must contend, and (iii) the crucial role of space-based observatories in efforts to characterize hot Jupiter planets. I also present the returns of an informal poll of the conference attendees conducted on the last day of the meeting, which may be amusing to revisit a decade hence.

1. Dominant Themes from the Meeting

Prior to the October 1995 announcement of the discovery of the planetary companion to 51 Peg (Mayor & Queloz 1995), few astronomers foresaw the existence of hot Jupiter exoplanets (for a notable exception, see Struve 1952). In the past decade, this subclass of exoplanets has been by far the best observationally studied, due to the rich set of follow-up techniques enabled by the proximity of the planet to the host star. The entire volume of this conference proceedings is replete with descriptions of dozens of exciting studies, some in progress, others planned for the near-future, detailing a wide diversity of observational efforts too numerous to list here. Radial velocity observations remain the dom-

inant tool for the detection of extrasolar planets, and provide precise estimates of the orbital parameters of these objects. However, these measurements alone yield little information about the planetary body directly (other than the value of the minimum mass). Once combined with complementary information, a rich, observationally-constrained picture of these objects emerges. An abbreviated list of such measurements include observations of the planetary photometric transits, reflected starlight, thermal emission, atmospheres and exospheres via transmission spectroscopy, as well as observations of the central star including its metallicity, magnetic activity, and the degree of alignment between the orbital and stellar rotational axes. While the breadth of observational and theoretical work presented at the meeting is far too great to be justly summarized here, I describe below several important themes that emerged over the course of the conference.

1.1 Impressive Leaps in Radial Velocity Precision

A significant development in the past couple years is that a radial velocity precision of $3 - 10 \text{ m s}^{-1}$, previously achieved by only a few teams, has now been put to wide-spread use by numerous groups around the globe. The benefit is that the additional researchers and telescopes now monitor a much more diverse set of primary stars than the F, G, K, and early-M dwarfs that have been the principal targets over the past decade (e.g. Marcy et al. 2005; Udry et al. 2006). Several examples of the parameter spaces under investigation are (1) targeted searches for hot Jupiters (Fischer et al. 2005; Bouchy et al. 2005), (2) surveys targeting only low-mass stars (Bonfils et al. 2005; Endl et al. 2003), (3) surveys dedicated to monitoring binary stars (Martínez Fiorenzano et al. 2005; Konacki 2005), (4) searches for planets orbiting young stars (Esposito et al. 2005), and, (5) searches for planets orbiting evolved stars (Hatzes et al. 2005; Sato et al. 2005a). These numerous efforts promise a rich catch of planetary systems that will flush out the full parameter space of planet formation. The forefront of Doppler precision is now well below the level of 3 m s^{-1} : The HARPS instrument has yielded a precision of 1 m s^{-1} (Santos et al. 2004), and data presented at the conference hinted that further improvements are close at hand (e.g. Mayor et al. 2005). A precision of tens of cm s^{-1} would enable the mass determination of terrestrial planets orbiting within the habitable zone of low-mass stars, a very exciting prospect indeed.

1.2 The Ongoing Value of Individual Objects

Despite the benefits of a statistical analysis of the hot Jupiter population as a whole, the detection of an individual planet of novel character can still significantly impact the field. Consider the following recent

examples: (1) The planet orbiting the star HD 149026 (Sato et al. 2005b) has a radius significantly less than Jupiter ($0.726 \pm 0.064 R_{\text{Jup}}$; Charbonneau et al. 2006), indicating the presence of a large core of solid material. This would seemingly prove that this planet formed through core accretion, as opposed to gravitational instability. (2) The planet orbiting one member of the stellar triple HD 188753 presents, in turn, a challenge to the core-accretion model, since the binary companion would have likely truncated the protoplanetary disk to a radius less than 1.3 AU, which lies interior to the snow line. Recent work (Pfahl 2005) suggests that this system may have formed in the dense environment of a stellar cluster. (3) Two new hot Neptunes (HD 4308b; Udry et al. 2005, and GJ 581b; Bonfils et al. 2005) were the welcome news with which the conference began. The presence of such objects and a handful of others (Butler et al. 2004; McArthur et al. 2004; Santos et al. 2004; Rivera et al. 2005) may hint at a large population of low-mass planetary companions in short-period orbits. The detection of the first transiting Neptune-mass planet is eagerly awaited. (4) Soon after the conference ended, the announcement of the discovery of the transiting planet HD 189733 b (Bouchy et al. 2005) left many observers scrambling to gather data both from the ground and space. This enthusiasm was motivated by both the proximity of the system ($d = 19$ pc, making it the closest known transiting exoplanet), and the favorable ratio of the planet's area to that of the star, both of which facilitate observations geared to detect the planet directly in either emitted or reflected light.

1.3 The Need for Dynamical Mass Estimates

The controversy over the mass determinations (and hence planetary status) for the recently imaged companions to 2MASS1207 (Chauvin et al. 2005) and GQ Lup (Neuhäuser et al. 2005) arises primarily from the lack of direct constraints on the theoretical emission models, upon which the mass estimates are based. The most reliable means to resolve this issue would be to locate similar objects in systems for which the masses may be determined dynamically, and thus provide the strict (and unforgiving) constraints on such models, as is currently ongoing for M-dwarfs (e.g. Ribas 2005). Indeed the greatest asset of transiting exoplanets is that the masses and radii may be determined robustly.

1.4 The Crucial Role of Space-Based Observatories

One of the most interesting aspects of hot Jupiter exoplanets is the set of opportunities that these objects afford for direct study. In that regard, it is very important to note the pivotal contributions from several space-based observatories, some of which were designed before such planets (and hence observations of these objects) could have been fore-

seen. Consider the numerous attempts to study the atmospheres and exospheres of transiting hot Jupiter planets: A host of ground-based efforts (e.g. Brown et al. 2002; Bundy & Marcy 2000; Deming et al. 2005a; Moutou et al. 2001, 2003; Narita et al. 2005; Winn et al. 2004) have yielded only upper limits (albeit very useful ones), whereas the detections have all come from the Hubble Space Telescope (Charbonneau et al. 2002; Vidal-Madjar et al. 2003). Similarly, upper limits from ground-based attempts to measure the thermal emission from such planets (e.g. Richardson et al. 2003a, 2003b) have recently been met by robust detections with the Spitzer Space Telescope (Charbonneau et al. 2005; Deming et al. 2005b; for a comparison to theoretical models, see Burrows et al. 2005; Fortney et al. 2005; Seager et al. 2005). And, despite remarkable ground-based photometry of transit curves (e.g. Moutou et al. 2004; Charbonneau et al. 2006; Holman et al. 2006), no such efforts have approached the exquisite results from both the STIS and ACS instruments aboard the Hubble Space Telescope (e.g. Brown et al. 2001, 2005). The MOST satellite (Walker et al. 2003) will either detect reflected light from one or more hot Jupiters, or place very stringent upper limits on the albedos (Walker et al. 2005), perhaps finally bringing respite to frustrated ground-based searches for this signal (Leigh et al. 2003a, 2003b; Collier-Cameron et al. 2002; Charbonneau et al. 1999).

2. A Poll of Conference Attendees

The rapid pace of successes over the past decade have inspired many fond remembrances of the prevailing wisdom prior to October 1995, some of which may be more accurate than others. In advance preparation for the conference celebrating the twentieth anniversary of the discovery of 51 Peg b (to be held in 2015), I put four questions to the attendees in the final session of the meeting, on the afternoon of August 25th, 2005. In these questions, the participants were asked to speculate as to the time scale and technical methods of significant future advances in the field of exoplanet science. The purpose of this survey was to record informally the opinions of the conference attendees as to these important questions, so that the replies could be revisited for both interest and enjoyment a decade hence. In order to avoid burdening the questions with unduly technical and lengthy definitions, the precise meaning of certain phrases (such as “habitable zone” and “extraterrestrial life”) was left somewhat vague: I note that this ambiguity pervades the more rigorous (and refereed) literature as well. The questions, and the tabulated responses, are presented in Figures 1 & 2.

A quick summary of the responses is as follows. Most participants felt that the transit method and/or the radial velocity method would

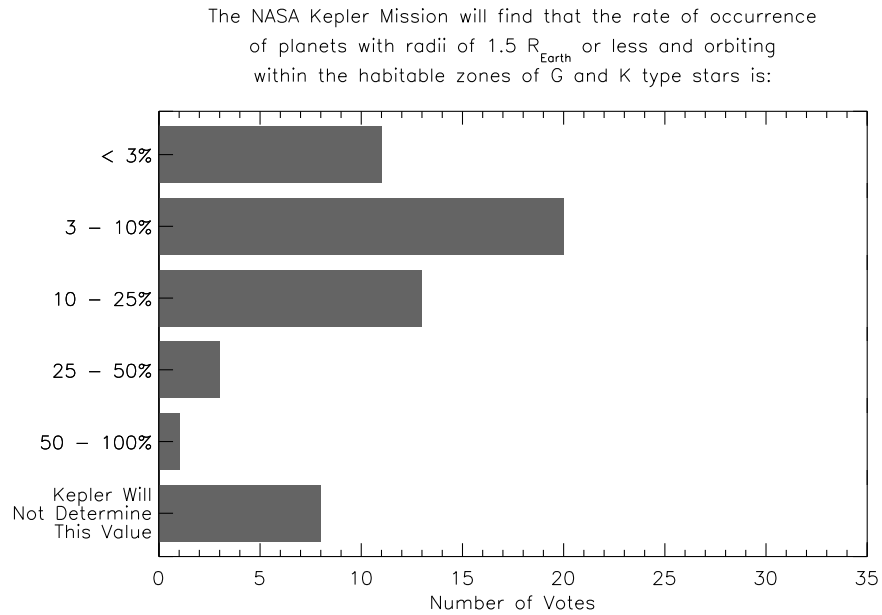
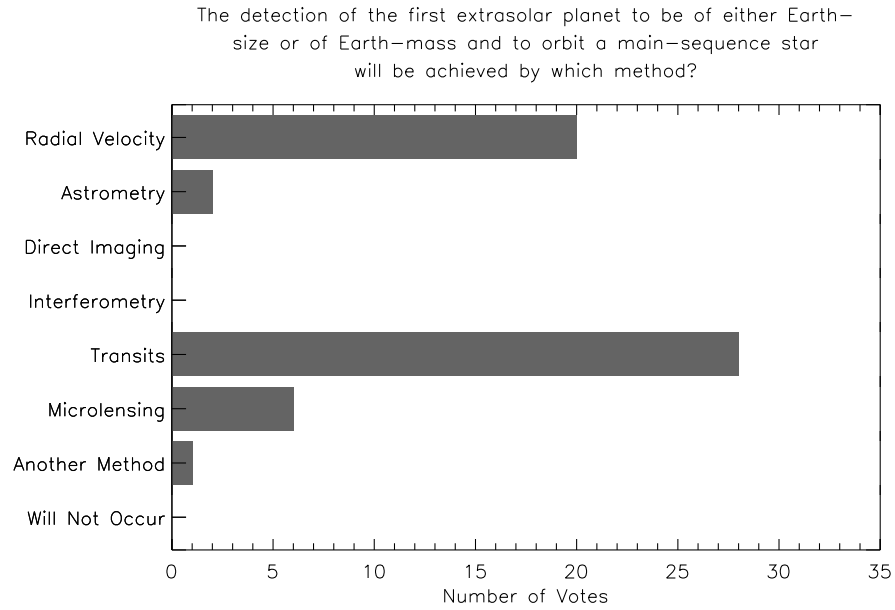


Figure 1.: Conference participants who attended the final session were asked to vote on four speculative questions (see Fig. 2). Upper panel: Attendees favored either the transit method or the Doppler technique to detect the first Earth-like planet. Lower panel: Attendees were divided as to the rate of occurrence of Earth-like planets as measured by the Kepler Mission, but few thought it would be greater than 25%.

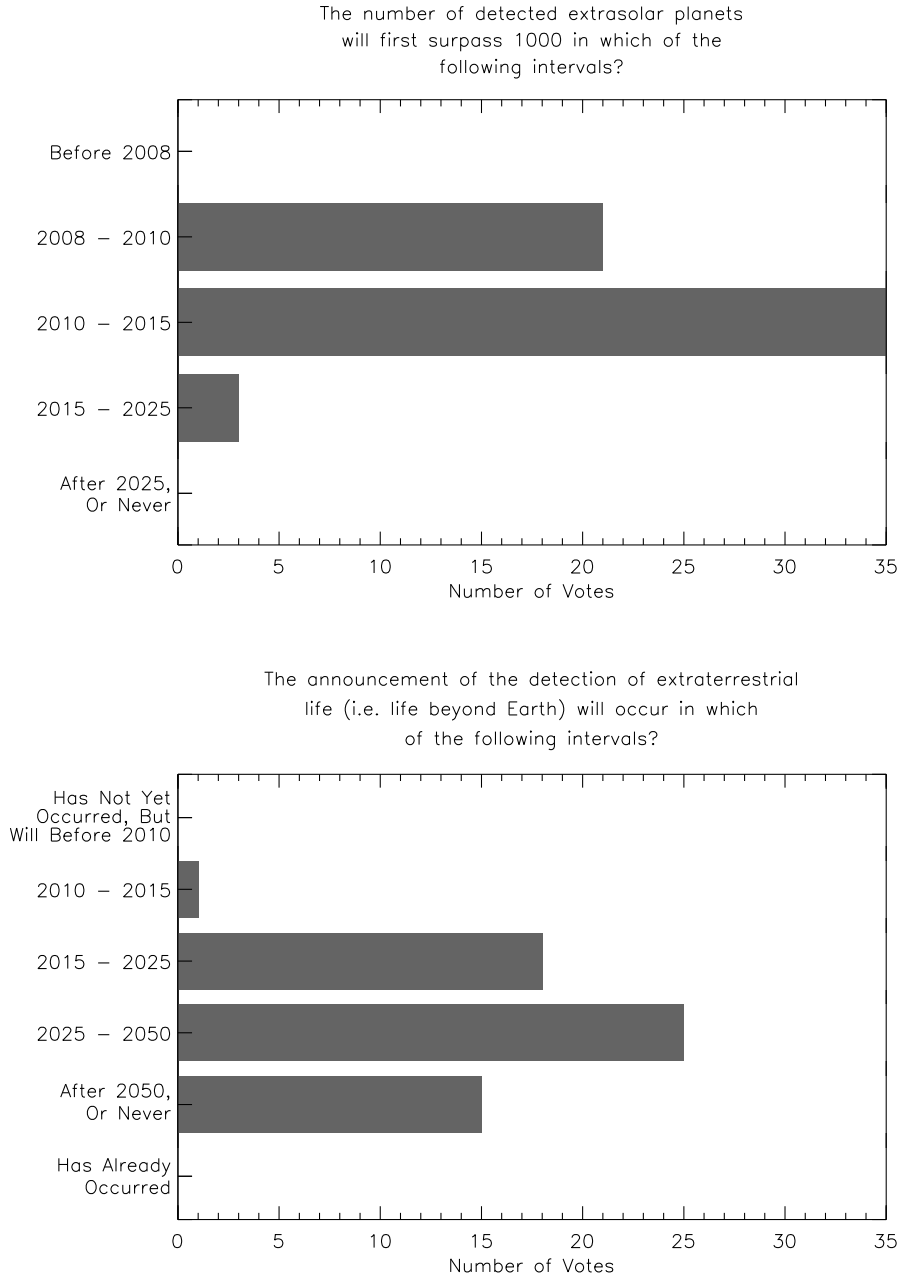


Figure 2.: Upper panel: *Roughly 160 extrasolar planets have been detected in the past decade. Most participants felt that rate of detection would be greater in the second decade, with 1000 such planets detected by 2015, 10 years hence.* Lower panel: *All but one of the attendees thought that extraterrestrial life would not be detected in the next decade, but a majority felt it would occur prior to 2050.*

yield the first detection of an Earth-like planet. Notably, no respondents felt that this would first be accomplished by interferometry or direct imaging (such as coronagraphy), efforts which currently receive a great deal of support. (It must be noted, of course, that missions such as TPF and Darwin are geared toward the spectral characterization of such planets, not simply their discovery.) Attendees were divided as to the value that the NASA Kepler Mission will determine for the rate of occurrence of Earth-like planets in the habitable zones of Sun-like stars. Few thought it would be greater than 25%, and a notable fraction questioned the mission's ability to determine this value at all. Whatever the method of discovery, nearly all participants felt that the number of detected exoplanets would exceed 1000 by the year 2015. As for the big question of extraterrestrial life, all but one of the attendees felt that its discovery would not occur by that date. A majority voted that its detection would be achieved prior to 2050, but a significant number felt it would occur after that year, or not at all.

With the exciting discoveries of the previous decade as our guide, we can only assume that the prevailing wisdom will, once again, be proven wholly unjustified.

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